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θ Canis Majoris ($\alpha = 6^h\ 50^m$; $\delta = -\ 11^\circ\ 55'$).

1897, December	15	$+\ 96^{km}$	CAMPBELL.
1899, October	16	$+\ 96.0$	WRIGHT.
1900, October	9	$+\ 95.5$	WRIGHT.

I *Pegasi* ($\alpha = 21^h\ 17^m$; $\delta = +\ 19^\circ\ 23'$).

1900, July	3	$-\ 75.7^{km}$	WRIGHT.
July	8	$-\ 74.9$	WRIGHT.
July	16	$-\ 77.1$	WRIGHT.

μ Sagittarii ($\alpha = 18^h\ 7^m.8$; $\delta = -\ 21^\circ\ 05'$).

1899, June	19	$-\ 75^{km}$	WRIGHT.
1900, May	30	$-\ 76$	WRIGHT.

These measures are subject to an uncertainty of several kilometers, on account of the character of the spectrum.

The negative sign indicates approach toward the solar system, and the positive sign recession from the solar system.

W. W. CAMPBELL.

NOTE ON THE PROBABLE ERROR OF MICROMETER MEASURES OF *EROS*.

During the recent *Eros* campaign for the determination of the value of the solar parallax, a great many observations were made visually with the micrometer. The question of the accuracy of these measures is therefore an important one. In general, the motion of *Eros* was so rapid in both coordinates that the probable error of the mean of a number of settings, even when made in rapid succession, could not be determined in the usual manner by treating the differences between the single settings and the mean as the *v*'s.

But on December 5, 1900, *Eros* ceased its apparent westward motion among the stars and began to move eastward. Consequently on that date its motion in Right Ascension was very slight — only about twenty-two seconds of arc in eight hours.

On the night mentioned, I measured the position of *Eros* with respect to two small stars in the early evening, when it was from three to two hours east of the meridian, and with respect to two other stars later in the night, when it was from four to five hours west of the meridian. In all I made eighteen measures of difference of Right Ascension, each consisting of ten settings with the micrometer. As the ten settings in no case required more than three and a half minutes, the motion of *Eros* could be entirely neglected and their variations from the mean treated as purely accidental errors of measurements.

In this way I found that the average probable error for the mean of ten settings was ± 0″.053, the smallest being ± 0″.036, and the largest 0″.085. During the evening measures the observing conditions were about the average. In the morning the seeing was poor. The average probable error for the evening observations was ± 0″.046, and for the morning ± 0″.058. The distances of the stars from *Eros* varied from 19″ to 133″, and their magnitudes from 9.1 to 12. In both evening and morning observations the two comparison-stars were on opposite sides of *Eros* in Right Ascension. Of the one hundred and eighty settings, eight give residuals from the mean of ten that exceed 0″.50, the largest being —0″.67.

As this use of the measures was not thought of until after they had been made, the results here given may be considered as fairly representing the accuracy, so far as accidental error of measurement is concerned, of my entire series of micrometer measures of *Eros*. R. G. AITKEN.

THE CROCKER ECLIPSE EXPEDITION TO SUMATRA.

The Crocker Eclipse Expedition to Sumatra sailed from San Francisco on February 19th, on the Nippon Maru. The expedition is in charge of Acting Astronomer C. D. PERRINE, who is accompanied by Assistant R. H. CURTISS. The Nippon reached Yokohama on March 11th, two days behind schedule time, unusually heavy weather having prevailed almost continuously. A cablegram from Mr. PERRINE was received at the Lick Observatory on April 5th, announcing the arrival of the expedition at Padang. The eclipse occurs May 17, 9^h 29^m P.M., Pacific standard time.

Weather permitting, it is hoped to secure observations with the following instruments:—

1. Five-inch aperture, 40-foot focus, Clark photoheliograph lens, for recording the details of the inner corona. The exposures will vary from ¼ second up to 2½ minutes, the latter on 18 × 22-inch plates.

2. Five-inch aperture, 67-inch focus, Floyd (Clark) photographic telescope, for securing a series of photographs showing the general features of the corona.